

Applying for Time at the Very Large Telescope Interferometer (VLTI)

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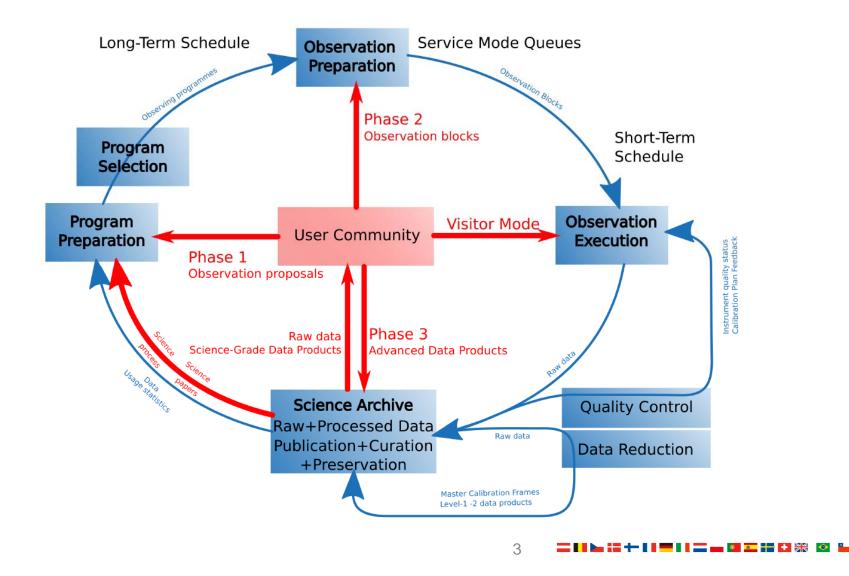




User Support Team









Applying for time

CfP released twice a year, March and September

General information and recent changes

> www.eso.org/sci/observing/phase1/p100/links.html

Technical details located in instrument web pages

> e.g.,www.eso.org/sci/facilities/paranal/instruments/gravity

Exposure time computation and constraints:

www.eso.org/observing/etc/ (!)

When feasibility of observation is established:

- Create/login ESO Portal account
- Download LaTex proposal package

Submit proposal before the deadline!



Non-member state proposals

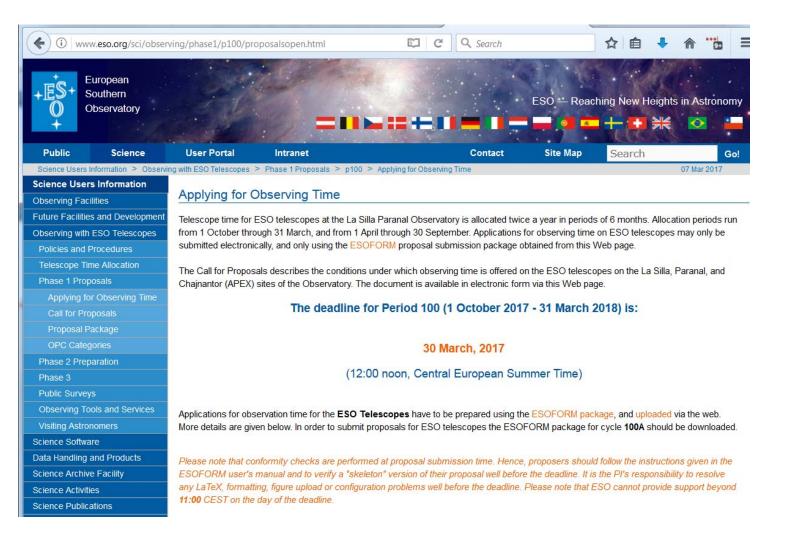
The policy for proposals presented by non-ESO-member state applicants is as follows:



- > a) The term "non-member state proposal" should apply by definition if at least 2/3 of the applicants are not affiliated to ESO member state institutes.
- > b) The following criteria will be adopted in evaluating non-member state proposals:
 - The proposal has to be scientifically outstanding.
 - The required telescope/instrumentation is not available at any other observatory accessible to the applicants.
 - If similar proposals of ESO members states and non-members state proposals are rated equally, preference will be given to the ESO member state proposals.
- Proposals for observing time may be submitted by scientists from any institution. However, ESO will only grant financial support to astronomers affiliated to institutions in the ESO member states.
 - See: www.eso.org/sci/observing/policies/Cou996-rev.pdf



Next deadline March 30, 2017





Data already available?

Science Users Information

Observing Facilities

Future Facilities and Development

Observing with ESO Telescopes

Policies and Procedures

Telescope Time Allocation

Phase 1 Proposals

Phase 2 Preparation

Phase 3

Phase 3 Overvie

Phase 3 Policies

Release Manager

Questions and Answers

News and Changes

Data Releases

Data Streams

Public Surveys

Observing Tools and Services

Visiting Astronomers

Phase 3

7 March 2017

 New release of imaging and multi-band catalogue data from the VST Public Survey ATLAS published [Read More]

6 March 2017

Data format specification for PIONIER Interferometric Data available [Read More]

20 January 2017

 Image and source list products released from the VISTA Variables in the Via Lactea Survey (VVV) [Read More]

[More News...]

In a nutshell, Phase 3 denotes the process of preparation, validation and ingestion of science data products (SDPs) for storage in the ESO science archive facility, and subsequent data publication to the scientific community. SDPs are produced by 1) principal investigators of ESO observing programmes, and 2) ESO pipelines as part of the quality control (QC) process or from specific, dedicated, re-processing projects for homogeneous raw data sets.

Phase 3 Quick Links

Phase 3 Main Level

ESO SDP Standard (v5) [PDF]

APEX Sub-mm Sky Maps [PDF]

Integral Field Spectroscopy: 3D Data Cubes [PDF]

Grantecan products: Spectra and Images [PDF]

Interferometric Data: PIONIER [PDF]

Phase 3 Release Manager

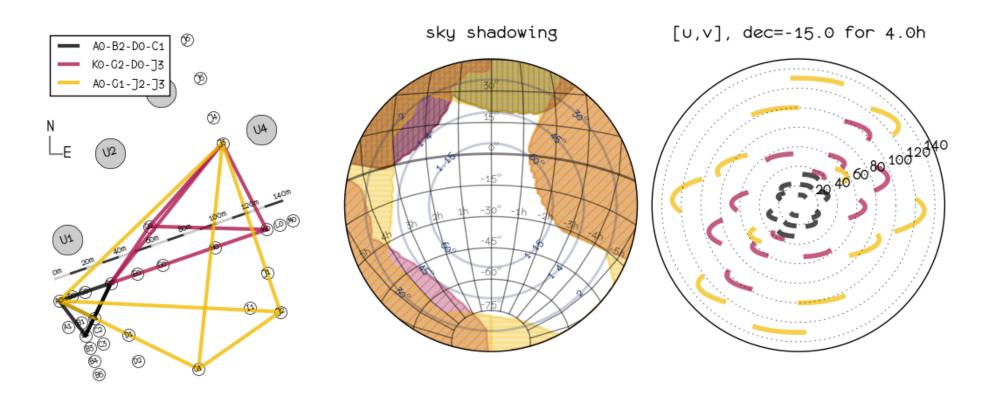
Get template for the data release description

Contact the Phase 3 Helpdesk



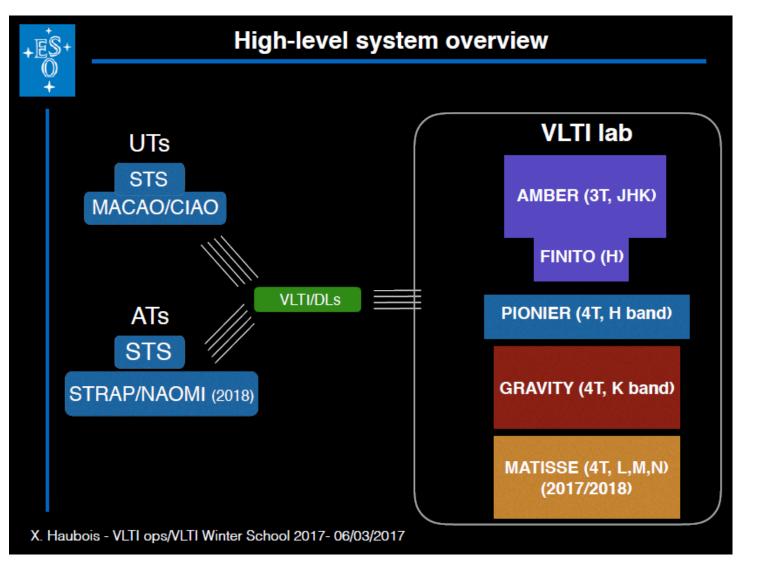
Available VLTI configurations

http://www.eso.org/sci/facilities/paranal/telescopes/vlti/configuration/P96.html.html





Instruments





Instrument specifications

	AMBER	PIONIER	GRAVITY	MATISSE
# of combined telescopes (ATs or UTs)	3	4	4	4
Spectral range and resolution	H-K (35,1500,12000)	H (none,30)	K (22,500,4000)	L,M,N (30-5000)
Fringe tracker	FINITO		Dedicated internal FT (on/ off-axis)	GRA4MAT
			 + astrometry offered in the near-future 	



Instrument limits (GRAVITY)

http://www.eso.org/sci/facilities/paranal/instruments/gravity/inst.html

	<0.6"	0.6"< seeing <1"	>1"
single field	7.0 ^m	6.0 ^m	5.0 ^m
FT dual field	7.5 ^m	6.5 ^m	5.5 ^m
SC dual field	7.5 ^m +3 ^m	6.5 ^m +3 ^m	5.5 ^m +3 ^m

Table 1. K-band limiting correlated magnitudes on ATs

Table 2. K-band limiting correlated magnitudes on UTs

	<0.6"	0.6" < seeing <1"	>1"
single field	10 ^m	9.0 ^m	8.0 ^m
FT dual field	10.5 ^m	9.5 ^m	8.5 ^m
SC dual field	10.5 ^m +3 ^m	9.5 ^m +3 ^m	8.5 ^m +3 ^m



Guaranteed time observations

Guaranteed Time Observations for Period 100

Please find below the protected target lists of the GTO teams for P100:

ARTEMIS

· List of protected observations for the ARTEMIS consortium

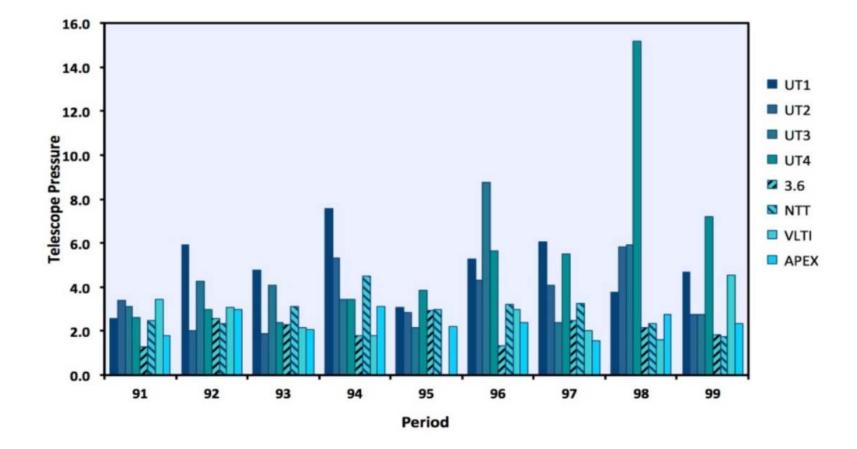
GRAVITY

· List of protected observations for the GRAVITY consortium

 List of protected observations for VISA-MPG 	Opening P100_GRAVITY-consortium.csv
KMOS • List of protected observations by the KMOS cor	You have chosen to open: P100_GRAVITY-consortium.csv which is: Microsoft Excel Comma Separated Values File (4,4 kB)
MUSE • List of protected observations for the MUSE cor	from: http://www.eso.org What should Firefox do with this file? Open with Microsoft Excel (default)
NACO List of protected observations for the PRIMA-DE 	 Save File Do this <u>a</u>utomatically for files like this from now on.
OMEGACAM List of protected observations for the INAF-OAC 	OK Cancel



Telescope presssure





What the OPC is looking for

Context and implications of the observations

- > What will we learn for a class of sources in general?
- > What questions will be answered, problems solved?
- Single targets need to be "high-profile"
 - > Are there going to be more general conclusions?

OPC appointed by ESO, but members must not be ESO staff

OPC comment for the programme:

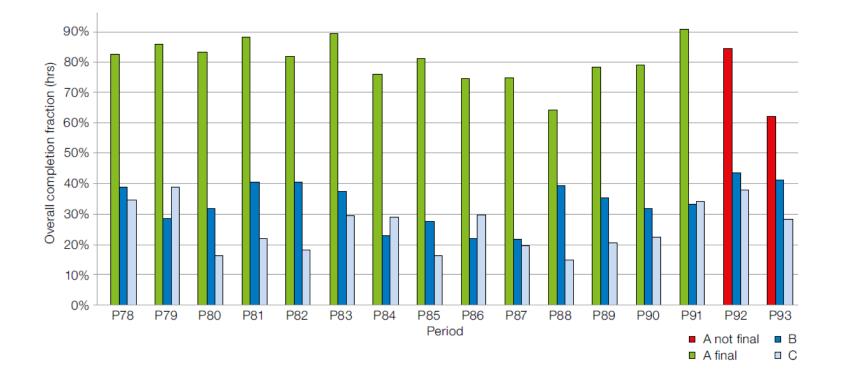
The panel agrees that VLT/GRAVITY observations of the innermost regions of the sgB[e] candidate in the Westerlund 1 cluster might well reveal important information about the influence of close binary interaction in this system. The proposal strengths lie in the characterization (geometry and kinematics) of the innermost part of the system with the goal to find a connection between the circumbinary torus and the outflow detected with ALMA. This will provide boundary conditions for the theory. However, the panel finds that the proposal does not present much information on the potential implications of results. Besides, the panel had the impression that the theoretical predictions on such connections still need to be constructed. A minor weakness perceived by the panel though still important, is the claim to reconstruct an image with only three hours of VLTI observing time. The panel deems image reconstruction with such a short exposure and hence sparsely filled UV plan unlikely.



OPC and scheduling

- OPC ranks proposals: A (highest), B, or C
 - > A: high-completion rate, carried-over if necessary
 - B: lower completion rate, impact on observational design
 - C: Filler, to be observable even in poor conditions
- Scheduling according to required constraints
 - Very good seeing: fewer of the highest ranked proposals
 - Same for photometric or clear conditions
- Visitor mode: only one chance!
 - Travel to Paranal ③
- Service mode: recommended for VLTI





F. Primas 2014, The Messenger 158



Observation preparation

Congratulations: you got time!

- Phase 2:
 - Prepare and submit observing blocks (P2PP tool)
 - OBs are verified before submission
 - Prepare short summary of observing instructions
 - Iterate, if necessary, with support astronomer (USD)
 - Pay attention to Phase 2 deadline
- Ask ESO user support in case of questions:
 - Send mail to <u>usd-help@eso.org</u>



Observation support

Subscribe to night reports

Receive notification when observations took place

Reply to issues reported by Paranal Operations:

- You would be contacted by your support astronomer (USD)
- Targets resolved?
- Calibrators bad?
- > Observations carried out on intermediate configuration



Operations

Queue scheduling based on rank index

- Select all OBs observable given the conditions
- Consider OPC rank and grade
- Consider time constraints:
 - Setting targets
 - Time intervals and monitoring
 - LST intervals

Planned improvements

- Cycle through each configuration once per month
 - Enable monitoring (important for young stellar objects)
- Make use of intermediate configurations
- > Automatic optimization of uv-coverage for imaging
- Selection of range of baseline lengths, not stations?